

# Sizes and shapes of spike sources observed by LOFAR

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# Introduction

- Low density and magnetic field in the upper corona of the Sun mean that low-frequency radio-emission is the only source of information about energetic particles in this part of the solar atmosphere
- Spatial distribution of coherent radio emission is a unique diagnostic tool for
  - propagating energetic particles
  - plasma density structure
  - plasma turbulencein the upper corona
- Electron acceleration and beam kinematics – e.g. *Carley et al. 2016, Reid & Kontar 2017, Morosan et al. 2019*  
Scattering and turbulence – e.g. *Kontar et al. 2017, Chrysaphi et al. 2018, Kontar et al. 2019, Gordovskyy et al. 2019, Chen et al. 2020, Murphy et al. 2021, Clarkson et al. 2023, Chen et al. 2023 ...*

# Spikes

- Clarkson et al. 2021 ApJL 917 L32
- Clarkson et al. 2023 ApJ 946 33

## First Frequency-time-resolved Imaging Spectroscopy Observations of Solar Radio Spikes

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## Solar Radio Spikes and Type IIIb Striae Manifestations of Subsecond Electron Acceleration Triggered by a Coronal Mass Ejection

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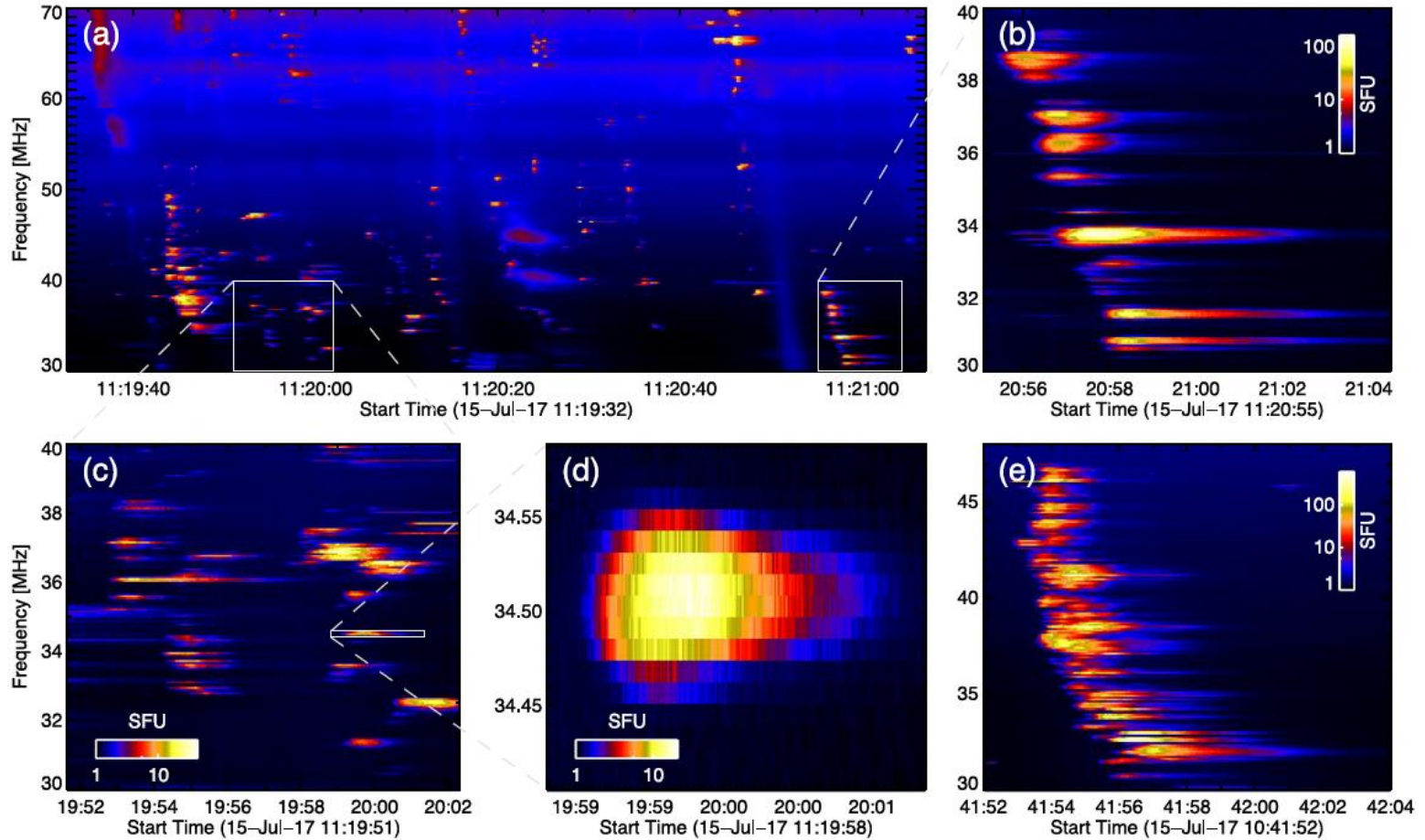
<sup>1</sup> School of Physics & Astronomy, University of Glasgow, Glasgow, G12 8QQ, UK

<sup>2</sup> LESIA, Observatoire de Paris, Université PSL, CNRS, Sorbonne Université, Université de Paris, 5 place Jules Janssen, F-92195 Meudon, France

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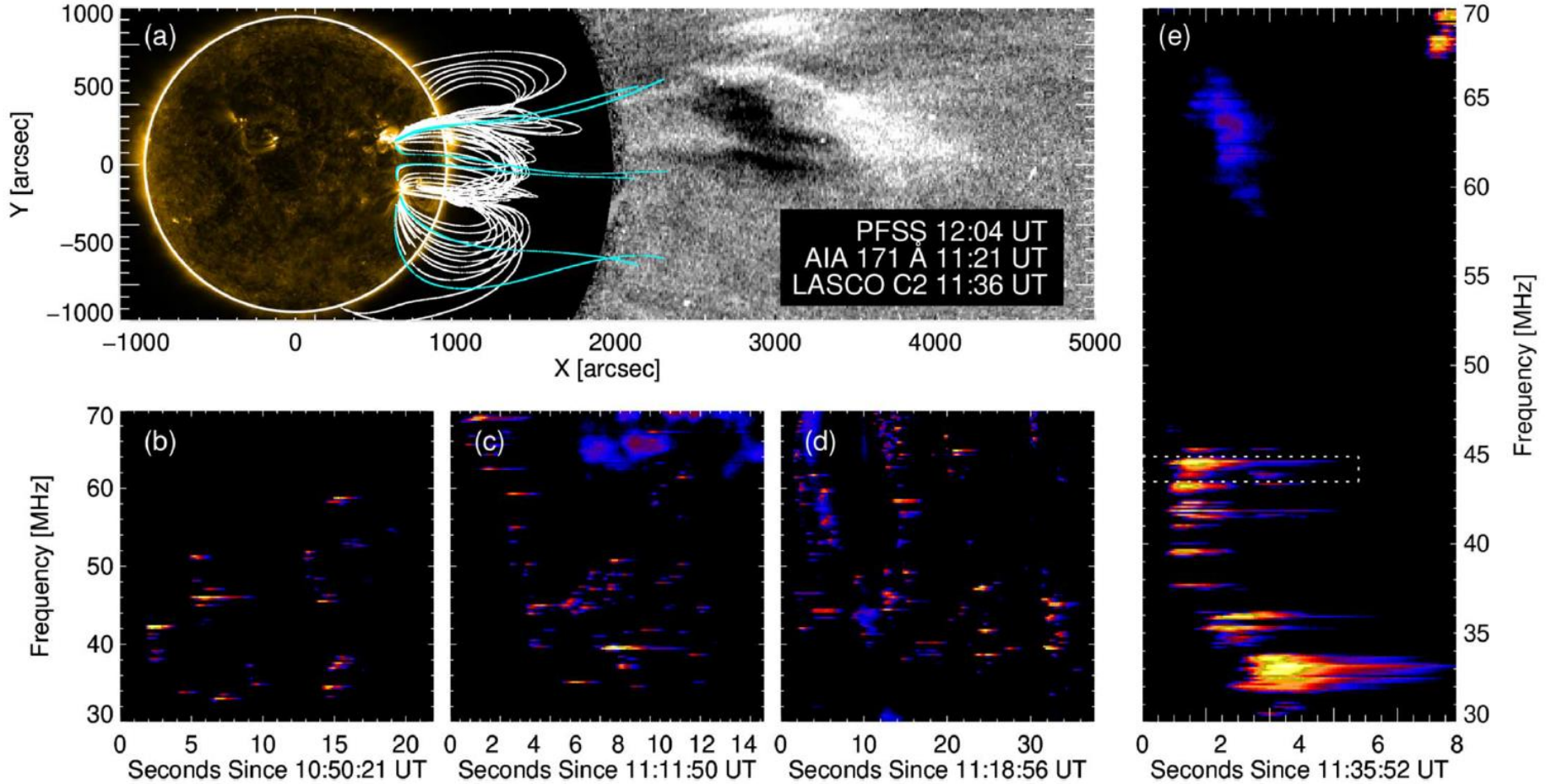
<sup>4</sup> Department of Physics & Astronomy, University of Manchester, Manchester M13 9PL, UK

# Spikes in DS



Clarkson et al. 2021 ApJL

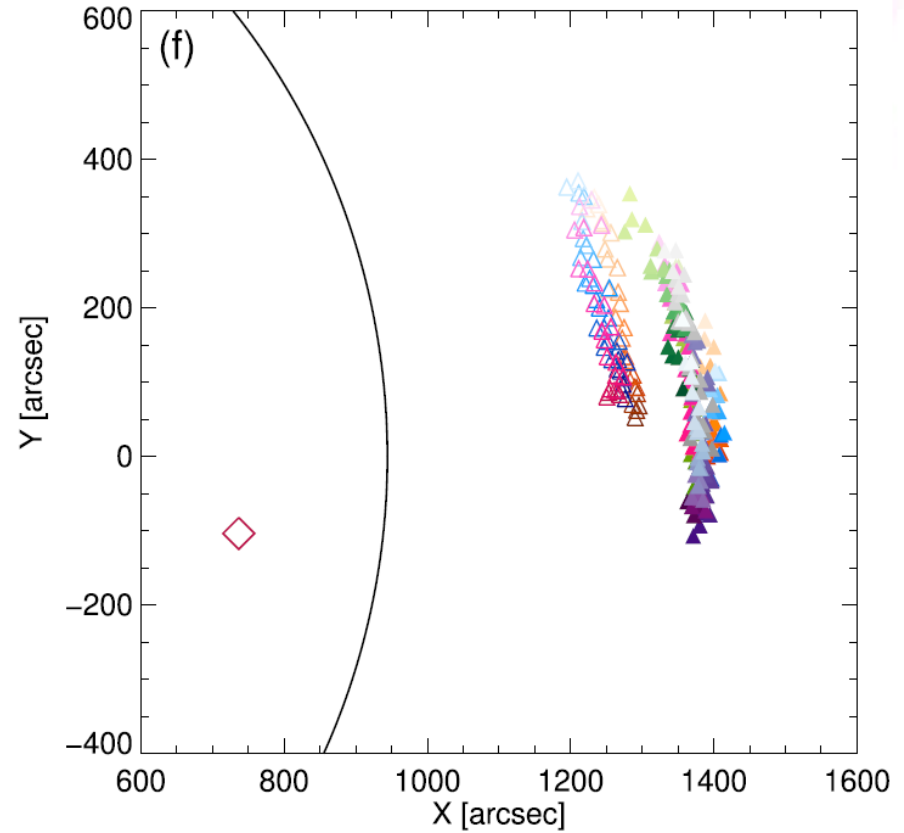
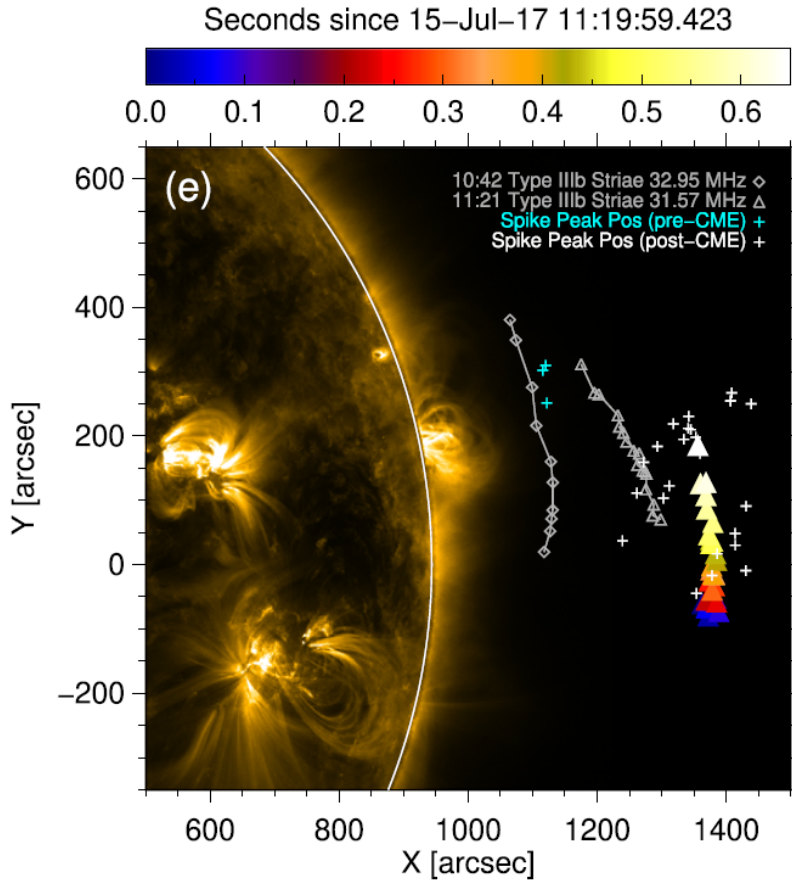
# Spikes in DS



Clarkson et al. 2023 ApJ

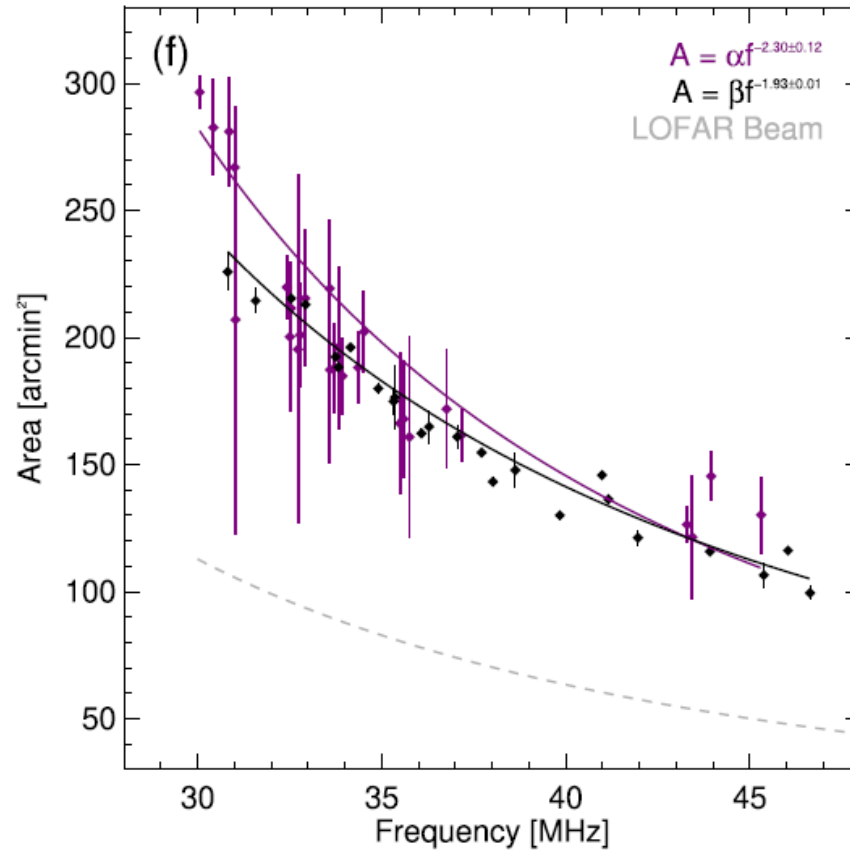


# Spikes – apparent motion of apparent sources



Clarkson et al. 2023 ApJ

# Spikes – apparent source sizes



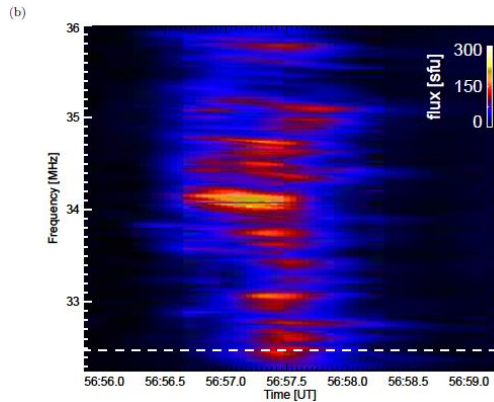
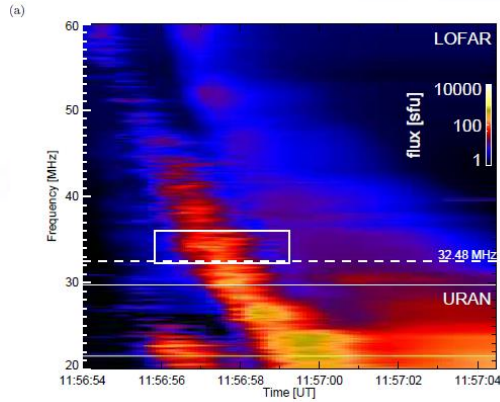
Clarkson et al. 2021 ApJL

# Spikes v “Normal” type III sources

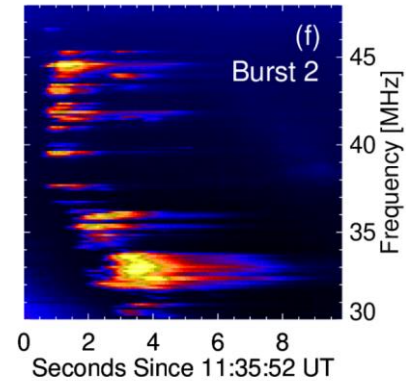
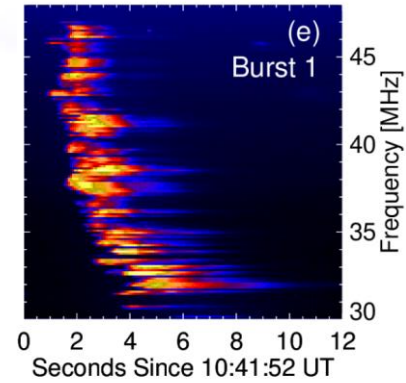
- Spikes are very compact as DS features compared to the striae
- Nearly co-spatial
- (Apparent) spike sources show very fast, sometimes superluminal, (apparent) motion
- Spikes appear due to an interplay of several factors, including short impulses of electron acceleration and strong, non-uniform scattering
- Their DS properties are different, but what about their geometry, i.e. sizes and shapes?



# Why not use the interferometric mode?



*from Kontar et al. 2017 NatCo*



*from Clarkson et al. 2022 in prep.*



**Solar observations require high temporal and spectral resolution (0.1s and 0.1MHz or better) and spatial resolution**



**TBA mode**

# Image deconvolution

$$I_d(f, A, z) = I_0(f, A, z - \Delta z_{\text{ir}}(f)) \star \mathcal{F}(f, A, z, \Delta A, \Delta z)$$

Observed image

Real image

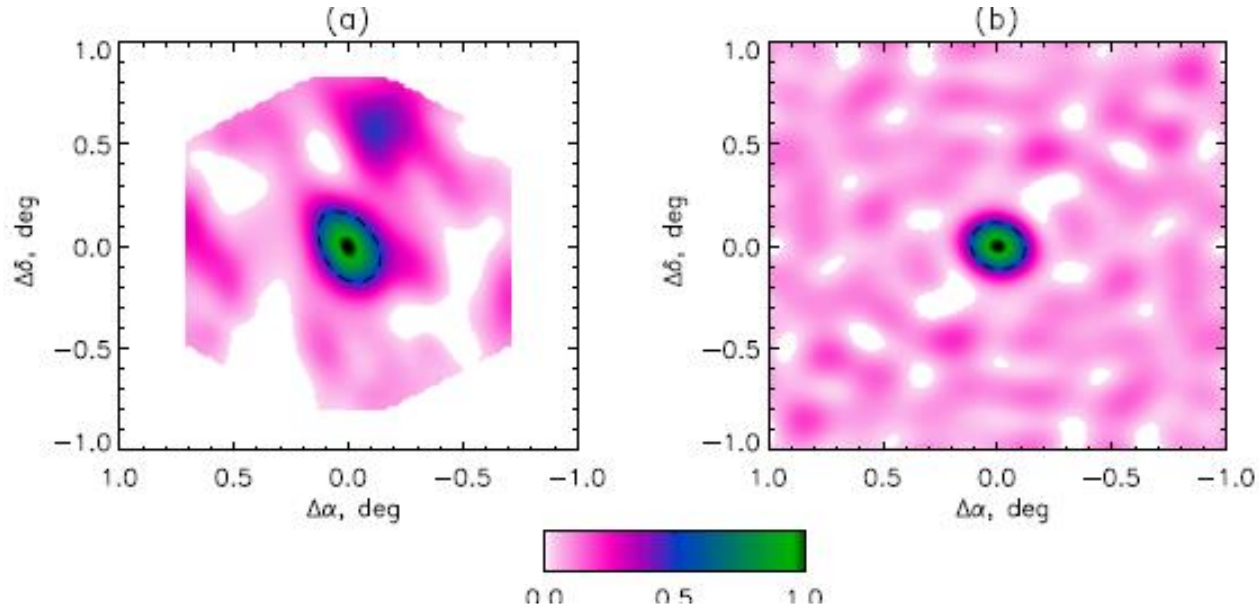
Point Spread Function PSF

Ionospheric refraction

## Do we know the PSF?

Tau A

Synthetic PSF

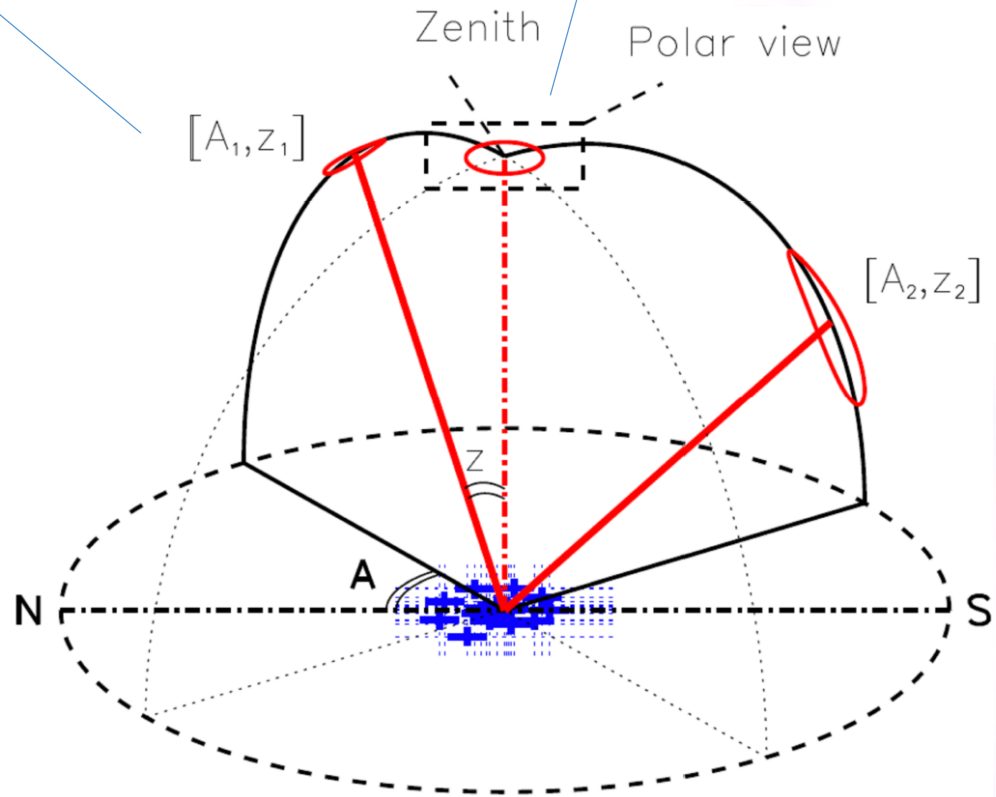


# PSF translation procedure

$$P(A=A_1, z=0) = f(\Delta A, \Delta z)$$

$$P(A=A_1, z=z_1) = f(\Delta A, \Delta z/\cos(z_1))$$

$$A=A_1 + \Delta A/\sin(z_1)$$



# PSF translation procedure

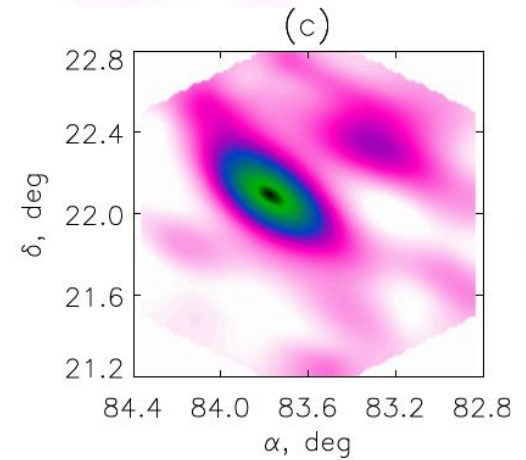
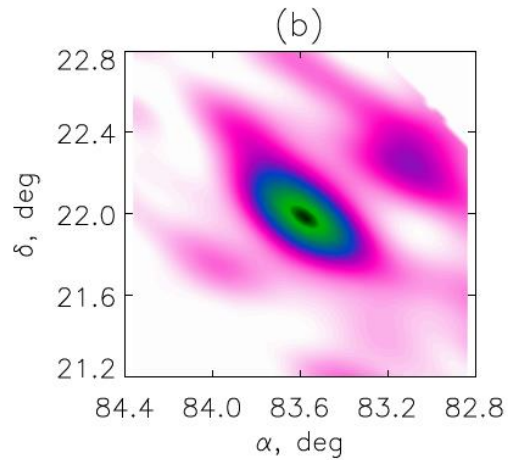
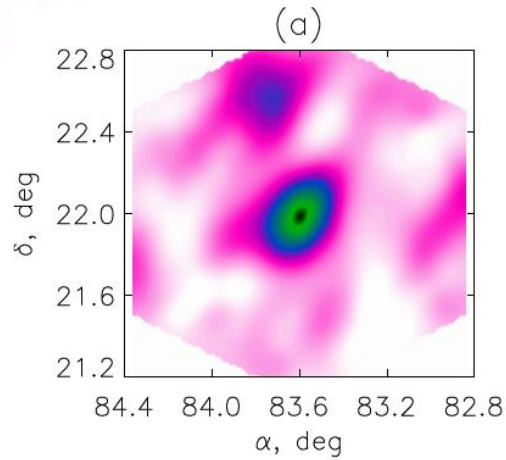
Tau A

Location 1

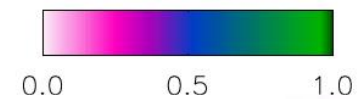
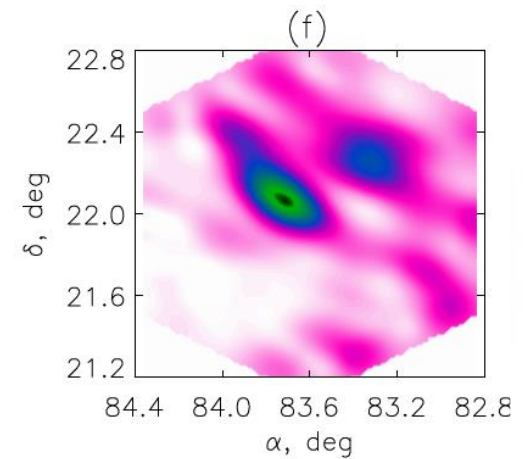
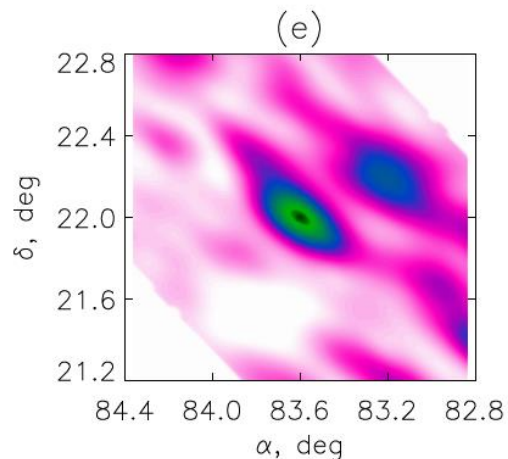
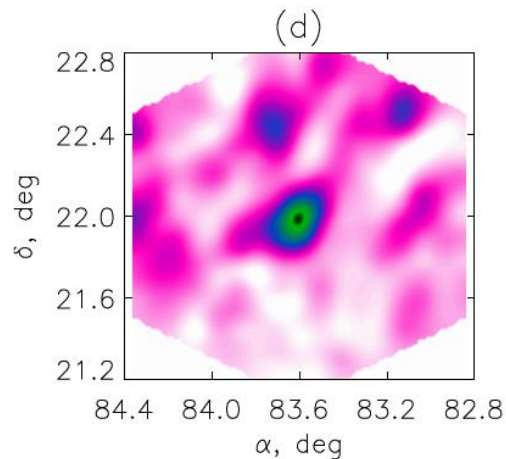
1 → 2

Location 2

30 MHz



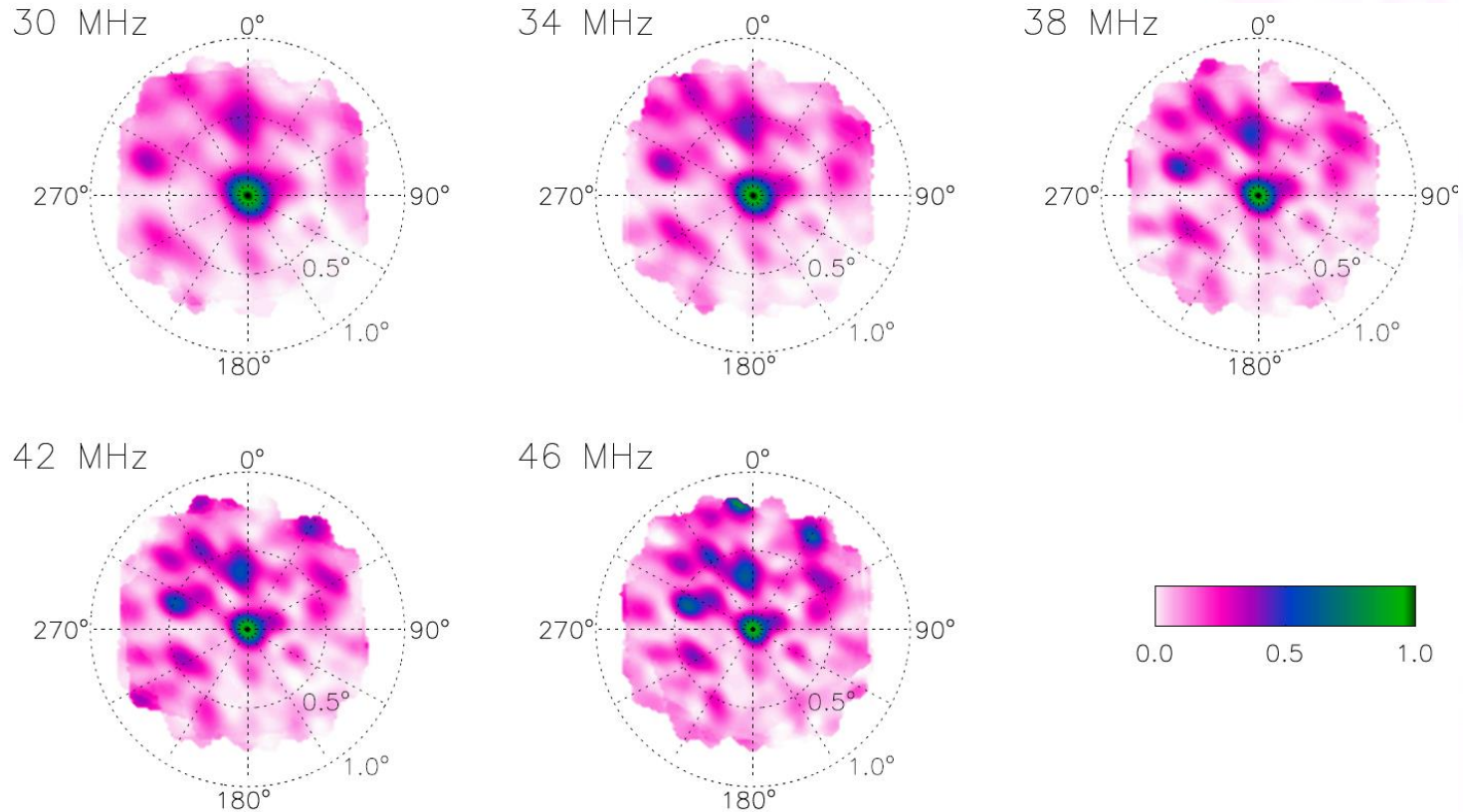
40 MHz



# LOFAR TBA PSF

- Around 30 individual observations of Tau A used
- Real Tau A shape is 'deconvolved out' of the observed intensity maps
- Intensity maps are 'translated' to zenith and combined

LOFAR TBA PSF at zenith



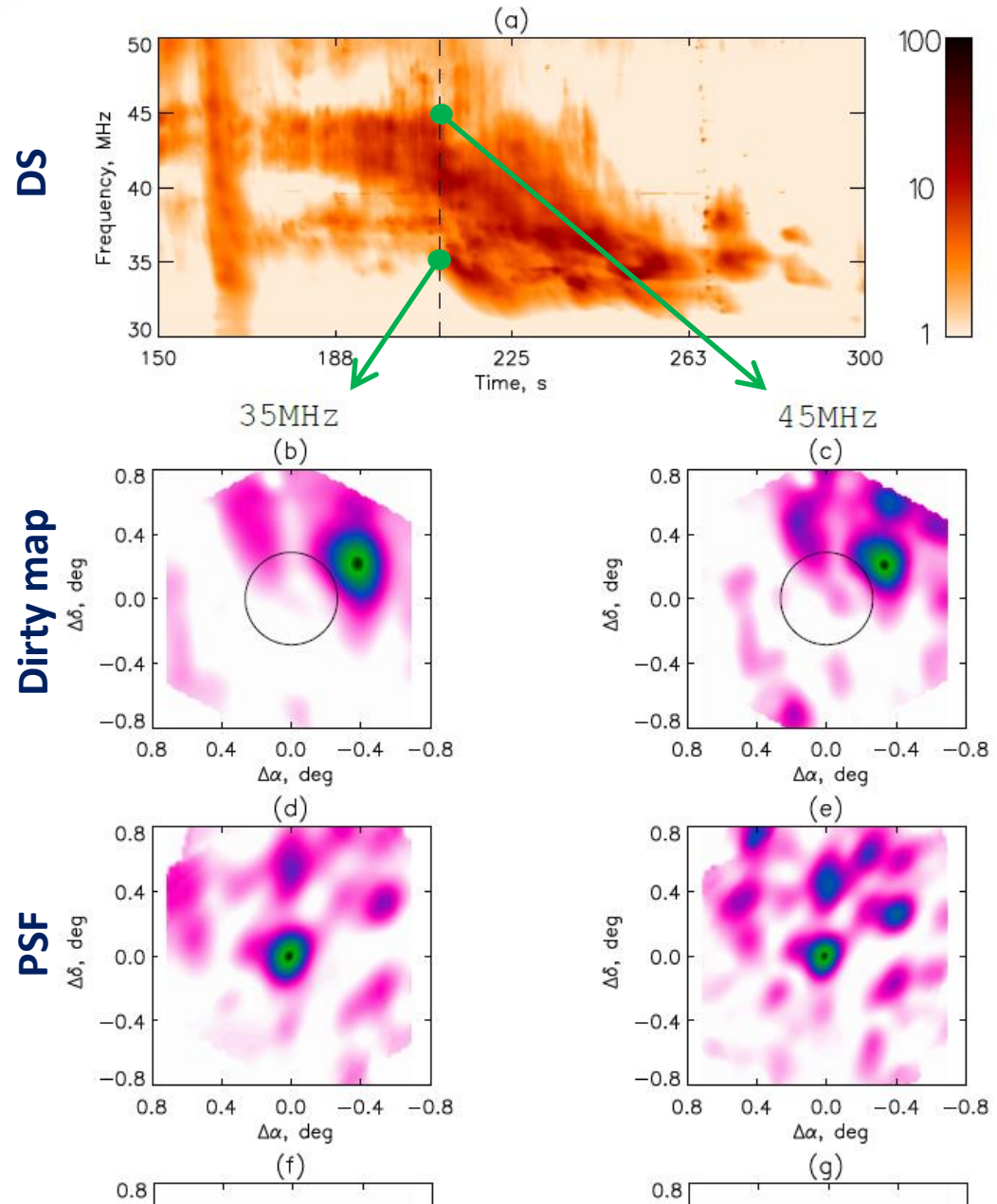


# Deconvolving solar images

1) Create intensity maps for selected positions at the dynamic spectrum

2) Produce PSFs for required position of the Sun and frequency

3) CLEAN

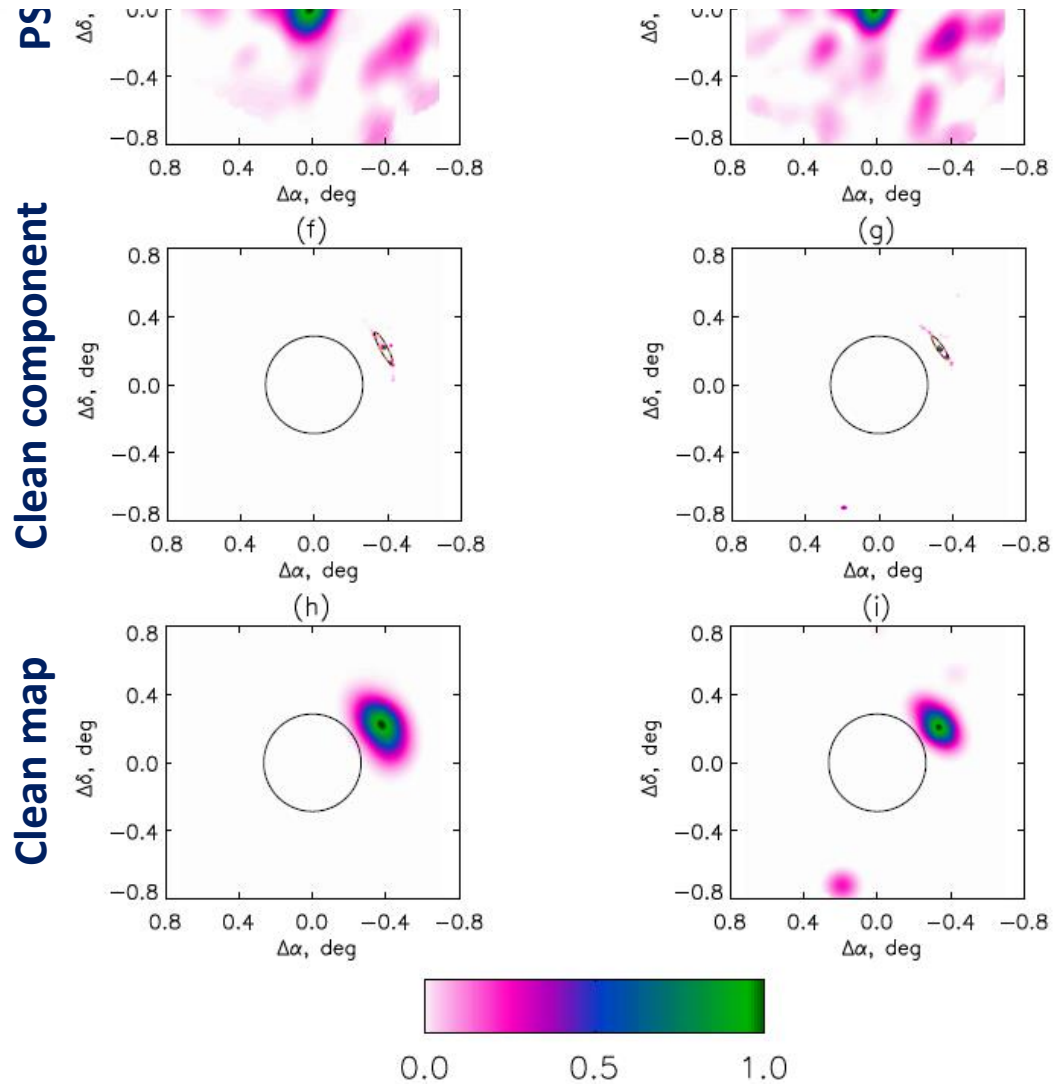




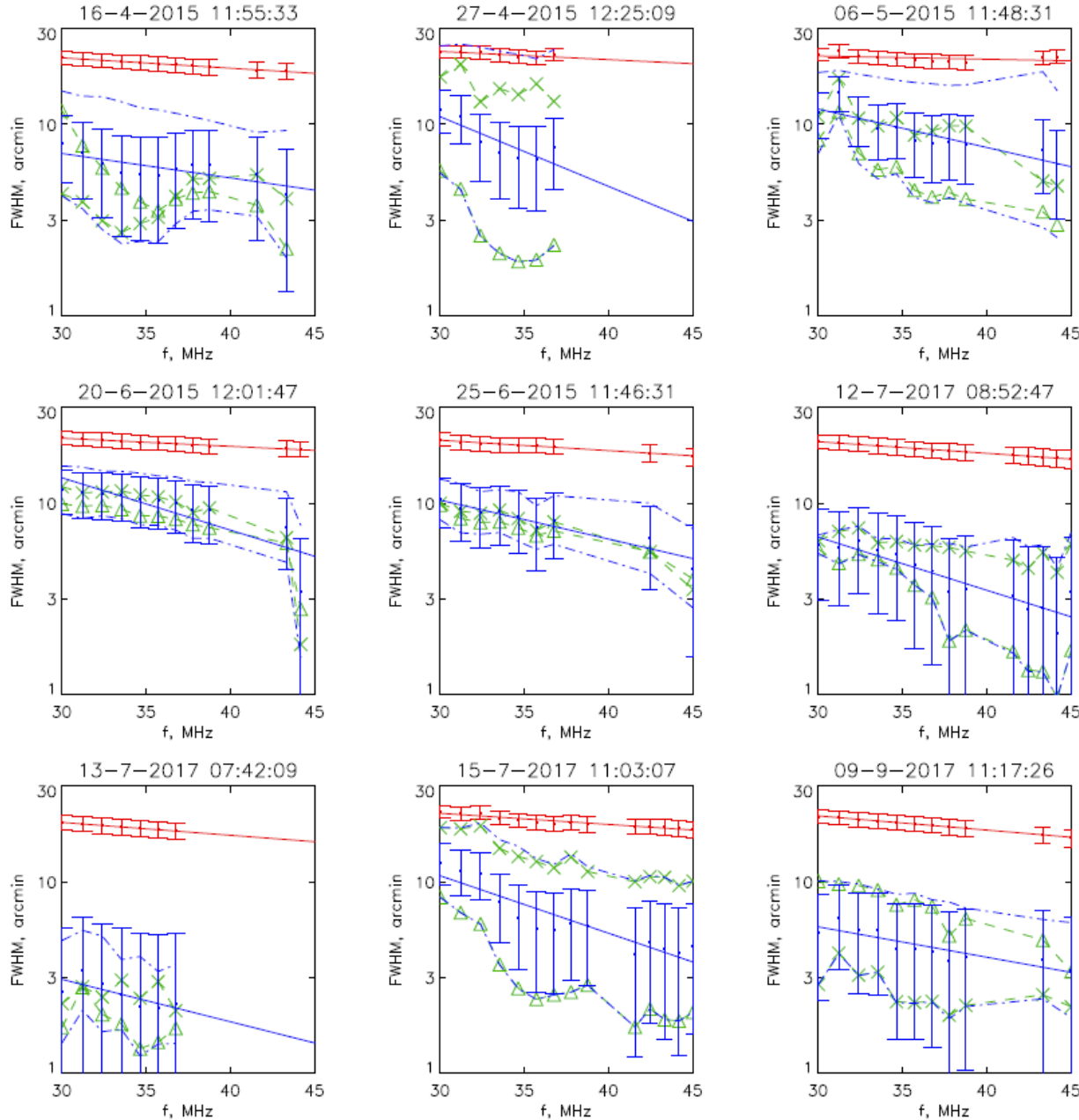
# Deconvolving solar images

4) Derive clean component maps and clean maps

5) 2D Gaussians with elliptical cross-sections to clean component maps



# Sizes and shapes of 'normal' type II-IV sources

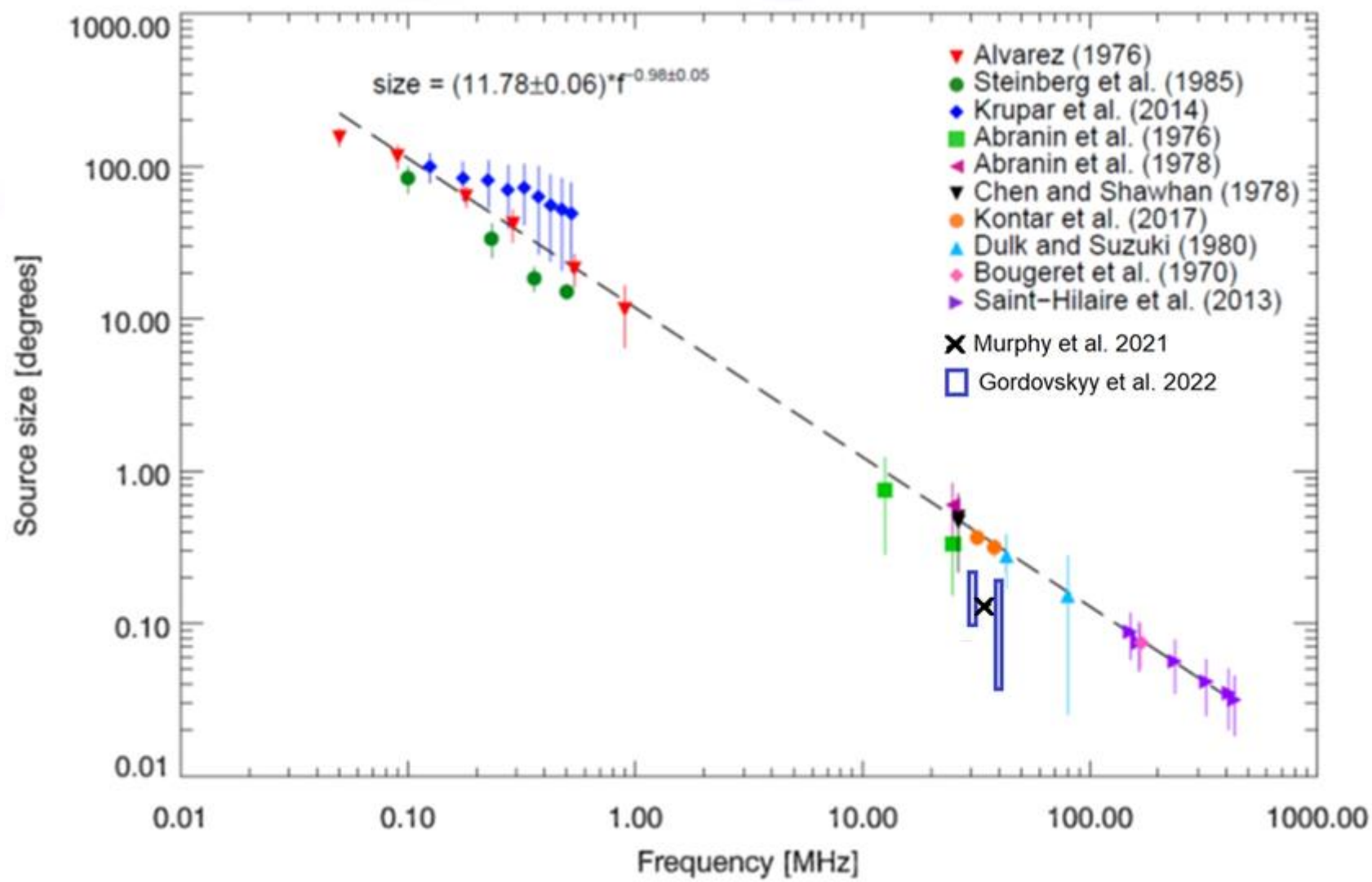


○ Typical sizes are 5-10 arcmin at 30MHz

○  $\sim f^{-1.8 \pm 0.5}$

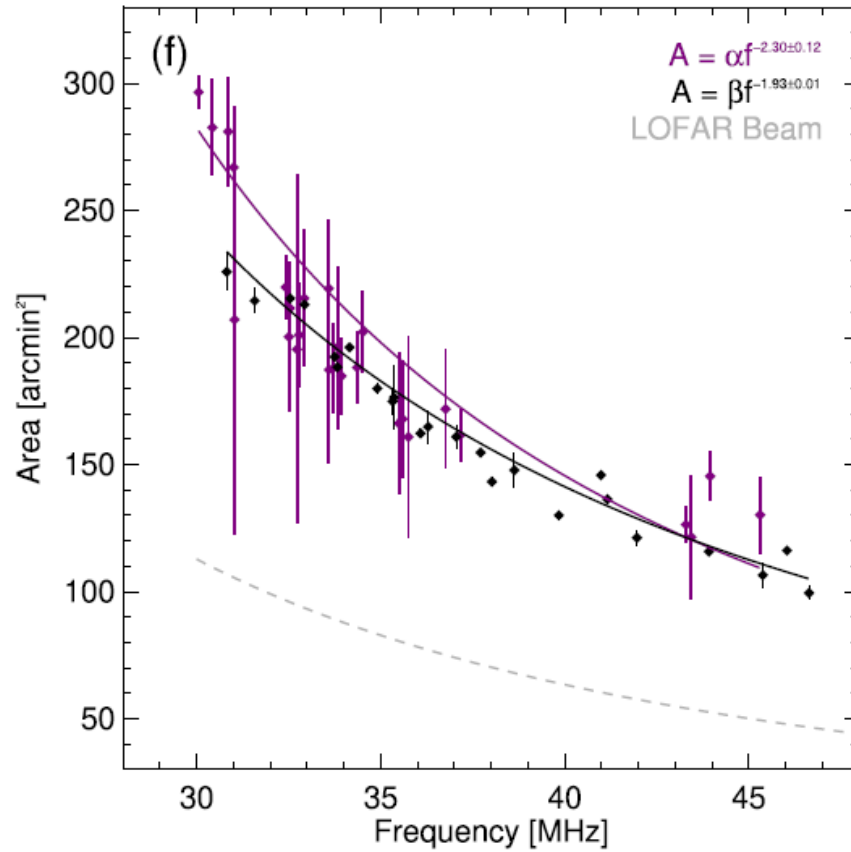
○ Small sample

# Sizes and shapes of 'normal' type II-IV sources



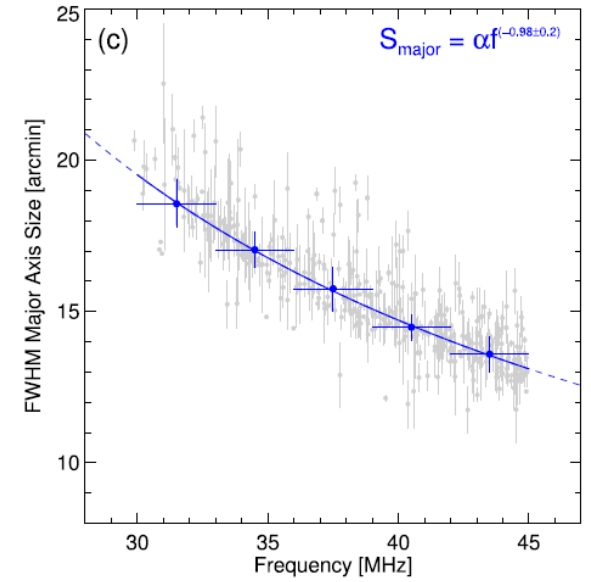
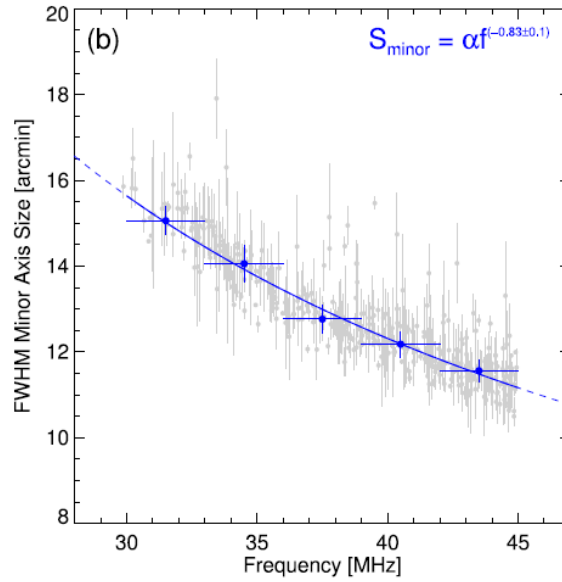
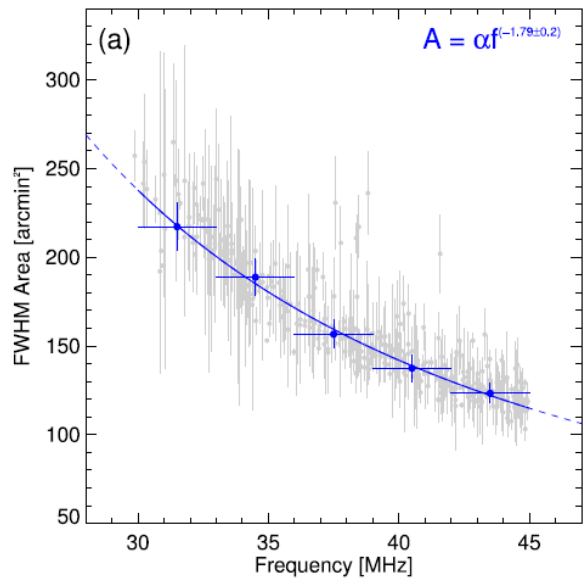
Background plot from Kontar et al. 2019

# Sizes of spike sources



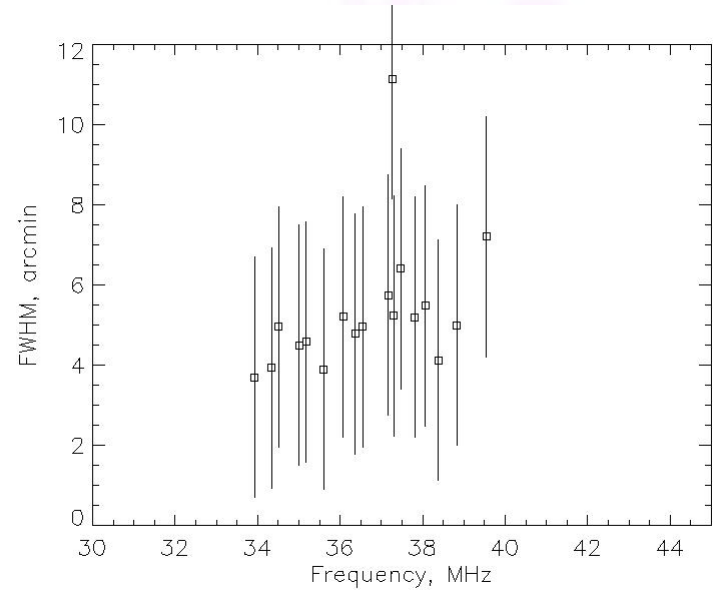
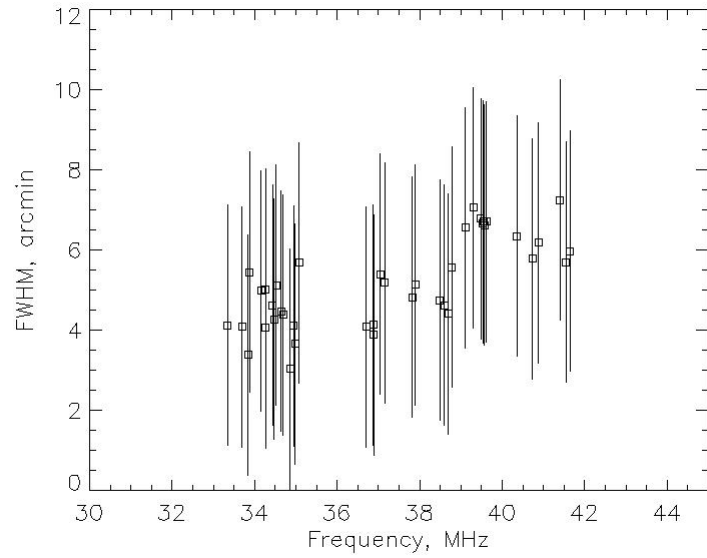
Clarkson et al. 2021 ApJL

# Sizes of spike sources



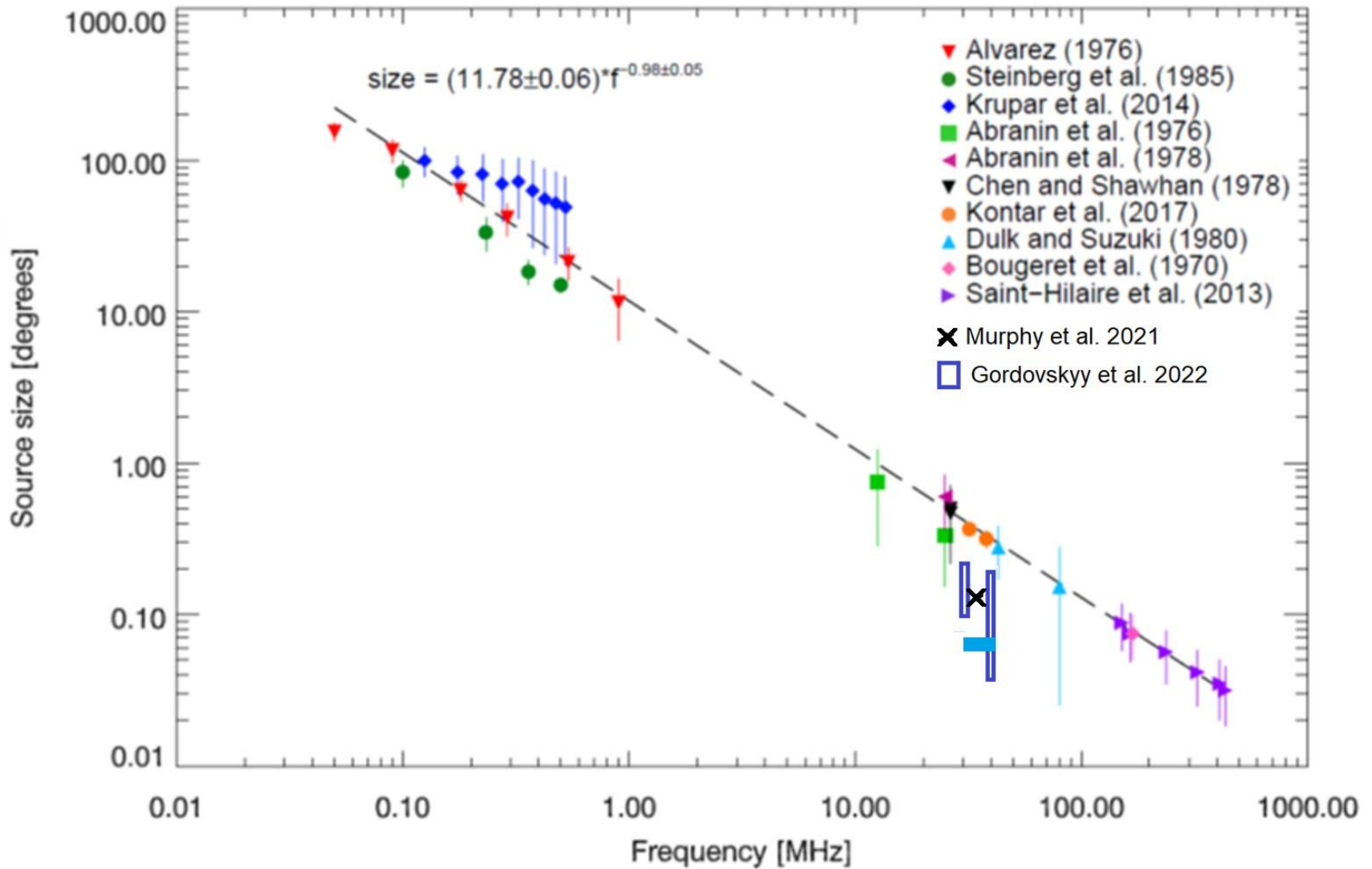
Clarkson et al. 2023 ApJ

# Sizes of spike sources



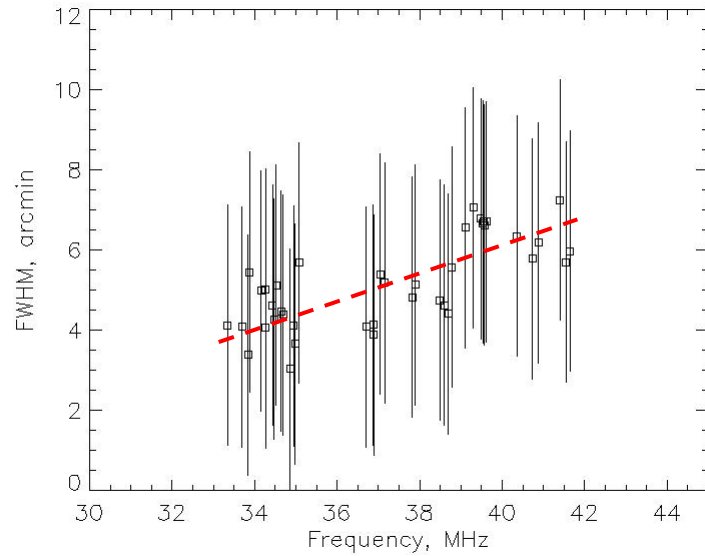


# Sizes of spike sources

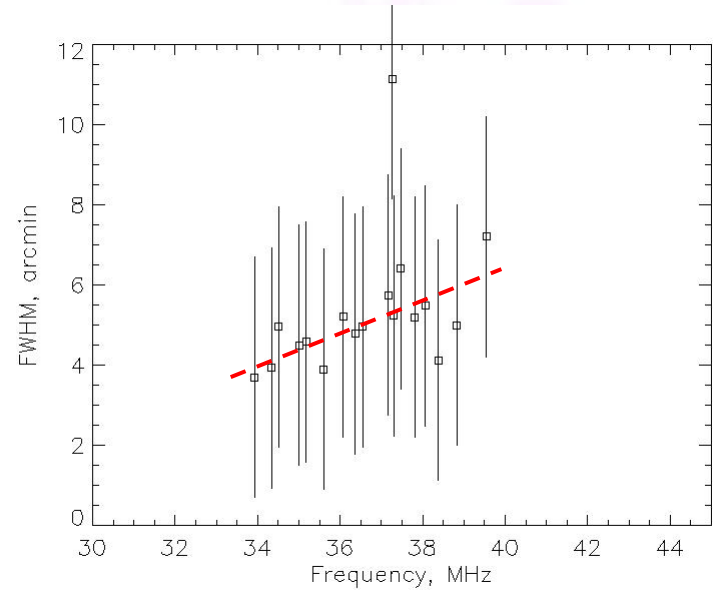


Background plot from Kontar et al. 2019

# Sizes of spike sources

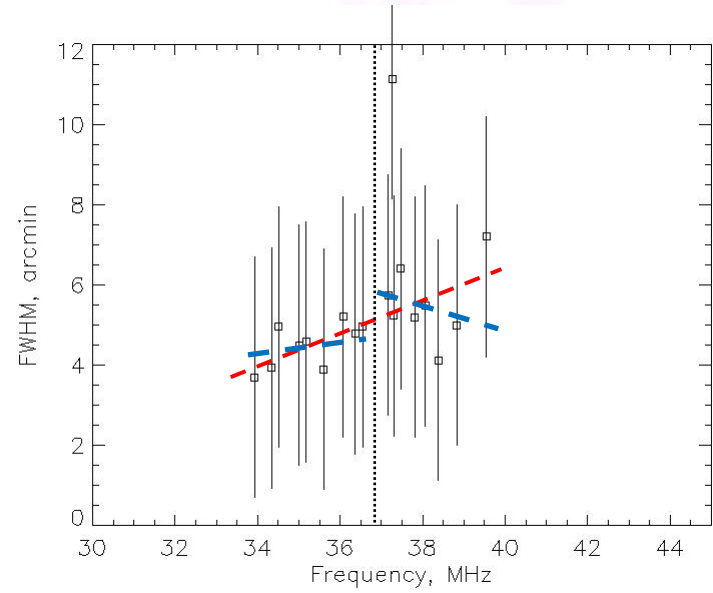
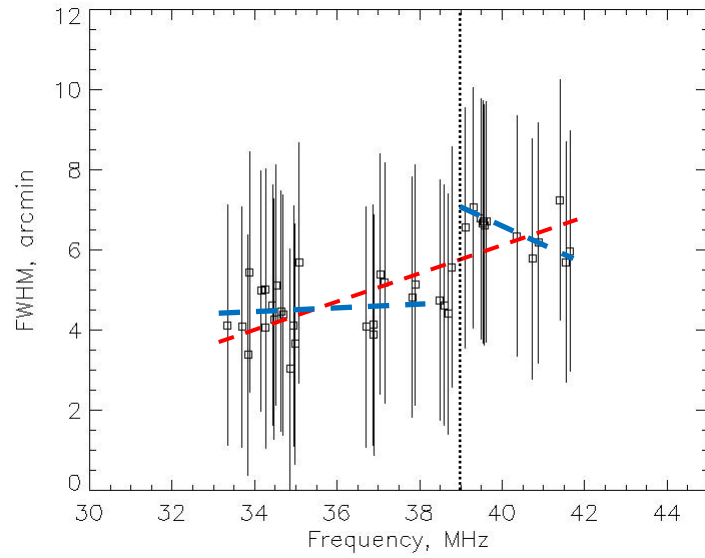


Power-law index 2.5

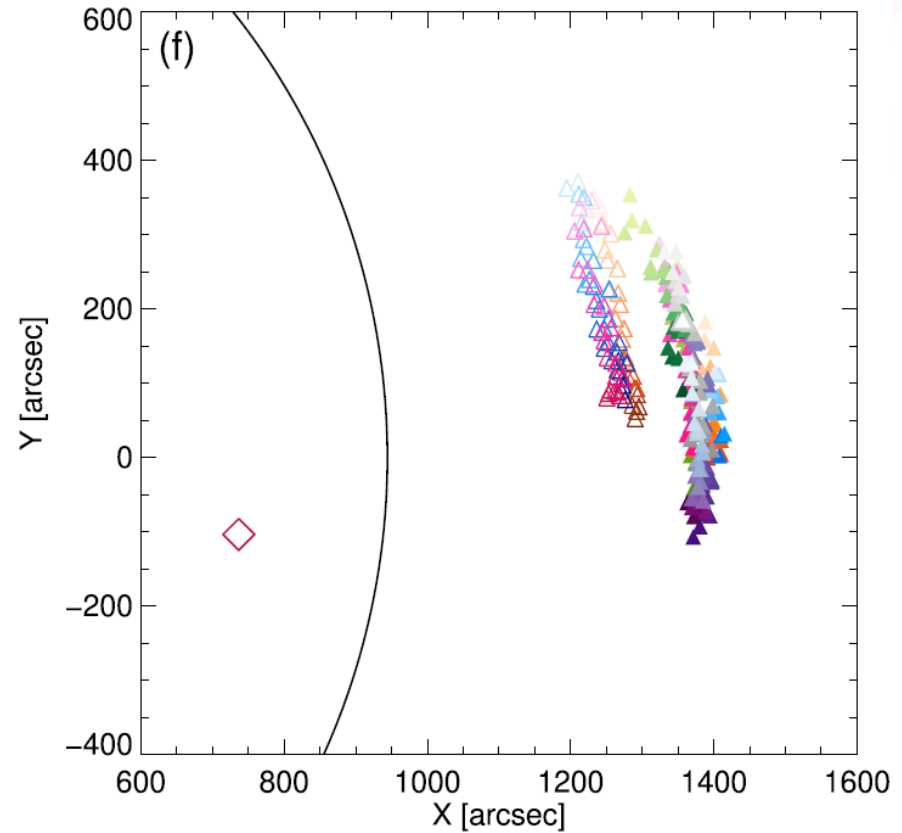
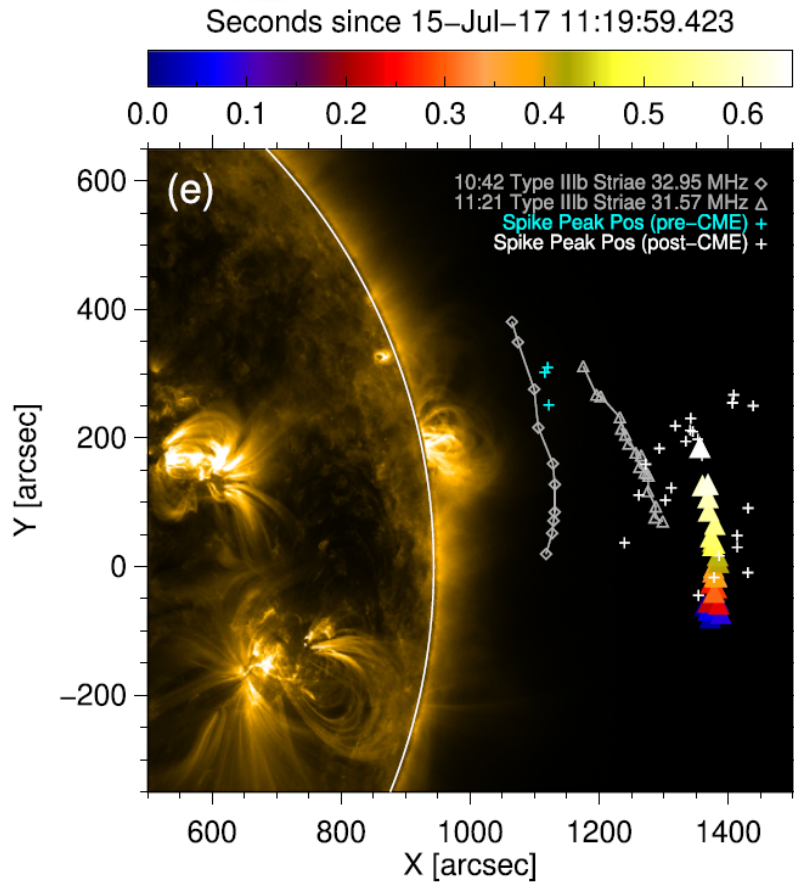


Power-law index 2.9

# Sizes of spike sources

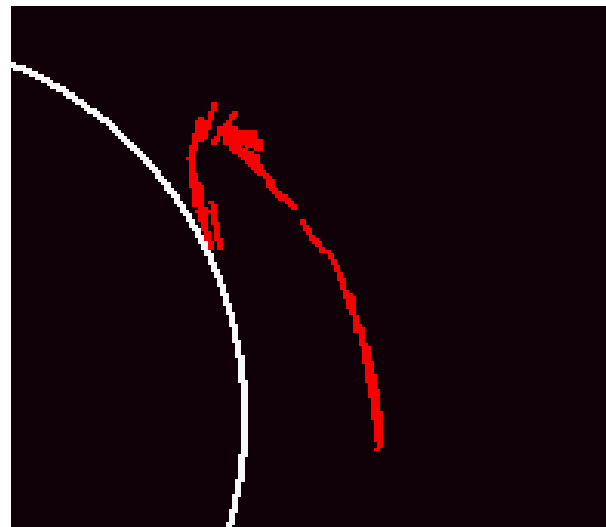
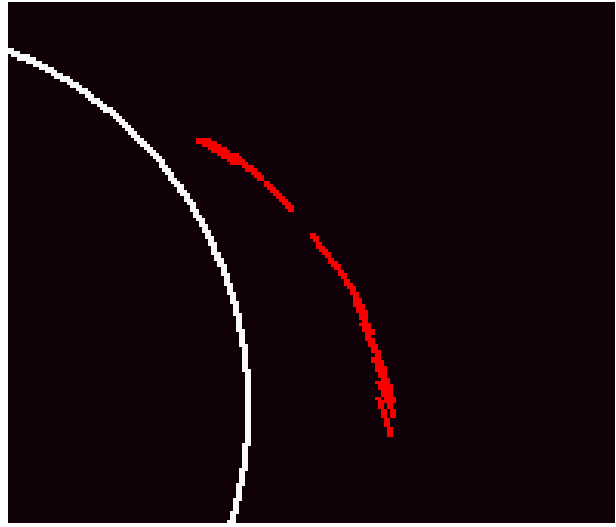
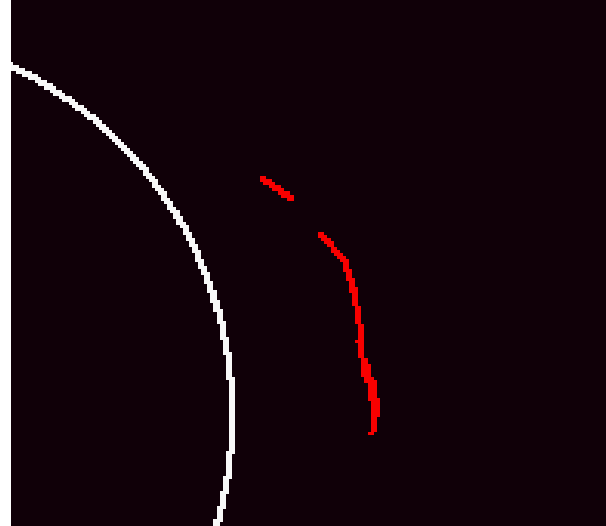
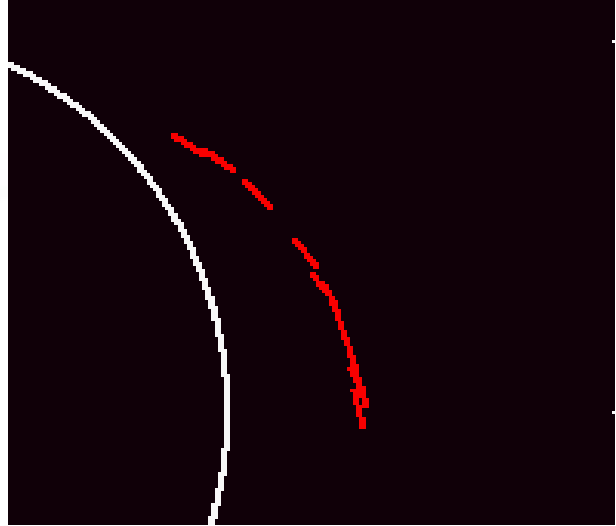


# Shapes and trajectories of spike sources

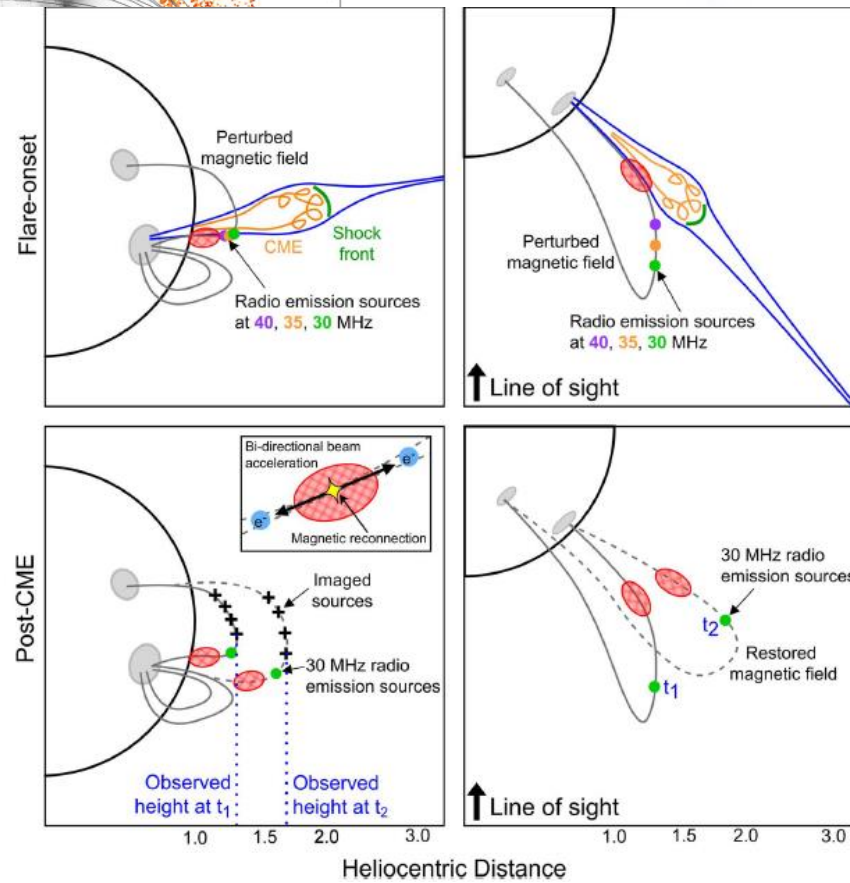
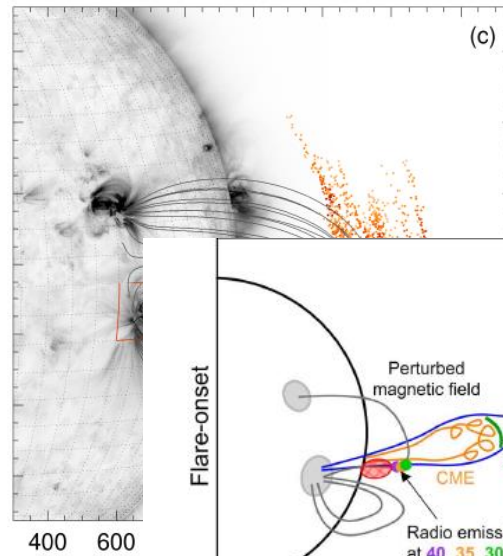
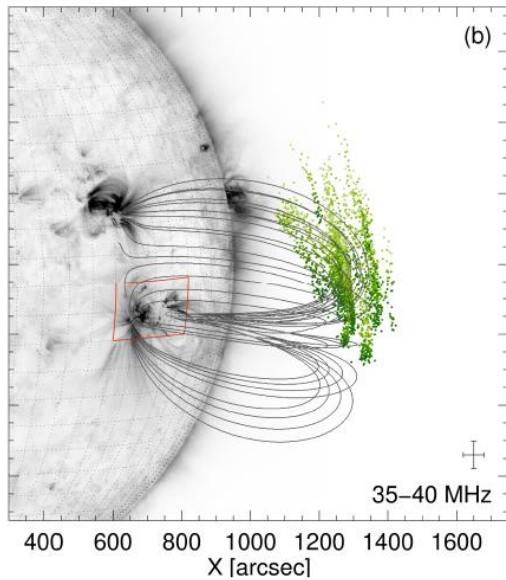


Clarkson et al. 2021 ApJ

# Shapes and trajectories of spike sources



# Shapes and trajectories of spike sources



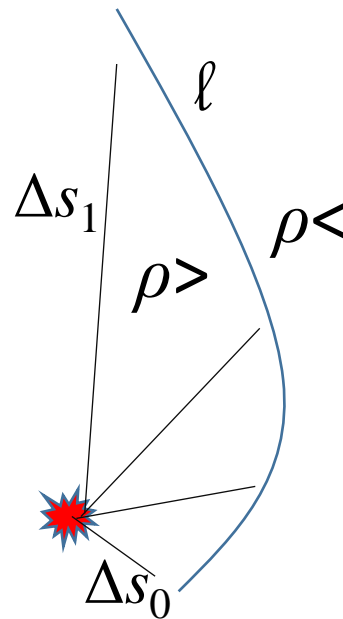


# Summary

- We have derived empirical PSF of LOFAR in TBA mode for the 30-50MHz range. It is very different from synthetic PSF. We can CLEAN solar images in arbitrary locations
- Found that spikes
  - **are typically 2-3 times smaller than “normal” type III sources**
  - show apparent motion with  $\sim c$ , **shape orientation correlates with the trajectory, and the trajectory is not  $r=\text{const}$**
- Phenomenological interpretation? – Same as in Clarkson et al. 2021, 2023, but **we need a shock**

## Why do we need a shock?

- It can help explain small sizes of spike sources
- Another possible explanation for the “trajectories” and superluminal velocities of apparent sources and



$$V_{\text{app}} = c \frac{\ell}{\Delta s_1 - \Delta s_0}$$